

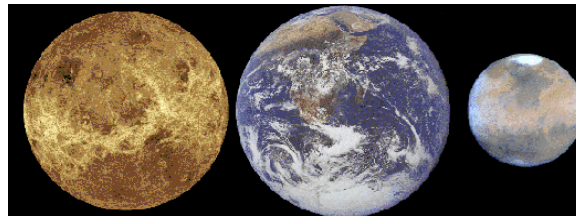
NON-LTE LIMB EMISSION AT 4.7 μM IN THE UPPER ATMOSPHERE OF VENUS, MARS AND EARTH: *OBSERVATIONS AND MODELLING*

**G.Gilli¹, M.A López-Valverde¹, B.Funke¹, M. López-Puertas¹
P.Drossart², G. Piccioni³**

¹Instituto de Astrofísica de Andalucía - CSIC, Granada, Spain

²Lesia, France

³IASF, Italy



Outline:

OBSERVATIONS

- CO limb observations @ 4.7 μm with Virtis/VEx, PFS/Mex, MIPAS/Envisat

THEORY AND MODEL

- Principal differences/similarities of terrestrial planets thermospheres
- Non-LTE model descriptions and results.

COMPARATIVE ANALYSIS

- Comparison model-data, planet-planet

SUMMARY AND CONCLUSIONS

Instruments characteristics and data set

VIRTIS/VEx: Visible Infra Red Thermal Imaging Spectrometer

VIRTIS-M = imaging spectrometer $R \approx 400$

V-M-vis 0.25 - 1 μm V-M-IR 1 - 5 μm IFOV $\approx 300\text{m/pix}$ periapsis

VIRTIS-H = spectrometer $R \approx 1800$ ($\Delta\nu \approx 1.3 \text{ cm}^{-1}$)

IR 2 - 5 μm IFOV $\approx 1.5 \text{ km/pix}$ periapsis

Day-time limb measurements from May 2006 to Oct 2008, about 460 Orbits.

PFS/Mex: Planetary Fourier Spectrometer

LW channel 5.5-45 μm (250 -1700 cm^{-1}) IFOV $\approx 7 \text{ km/pix}$ periapsis

SW channel 1.2-5.5 μm (1700-9200 cm^{-1}) IFOV $\approx 12 \text{ km/pix}$ periapsis

$\Delta\nu \approx 1.5 \text{ cm}^{-1}$

Two special orbits studied in the literature.

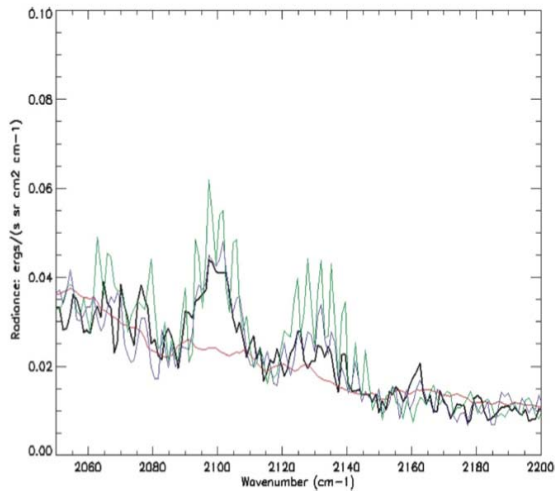
MIPAS/Envisat: Michelson Interferometer for Passive Atmospheric sounding

$\Delta\nu \approx 0.035 \text{ cm}^{-1}$ (unapodized) IFOV $\approx 3 \text{ km}$

Limb tracking observations 19-27 Sept 2002 = more than 6000 limb scans in the upper atmosphere

CO non-LTE limb observations on terrestrial planets:

PFS Mars observations:



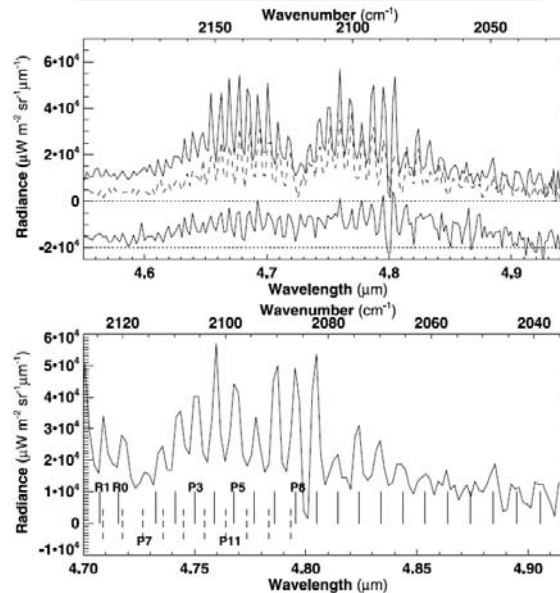
Formisano et al. 2006

Average spectra:

@ 80km-120km (green, blue)

@ 110km-180km (black)

VIRTIS Venus observations:



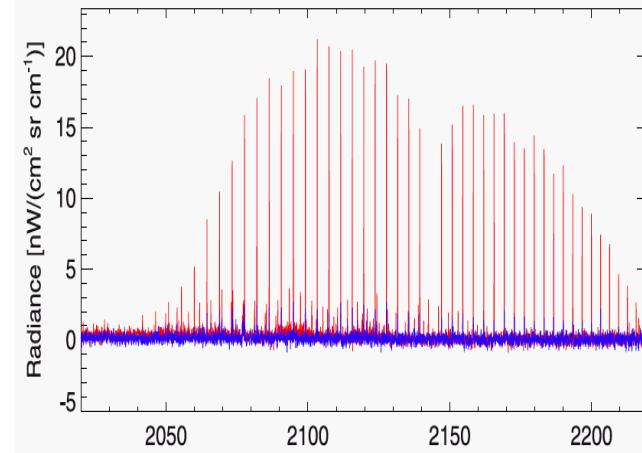
Gilli et al. 2009

Individual spectra:

@ 85 km, SZA= 60°, 70°

@ 118, SZA = 60°

MIPAS Earth observations:



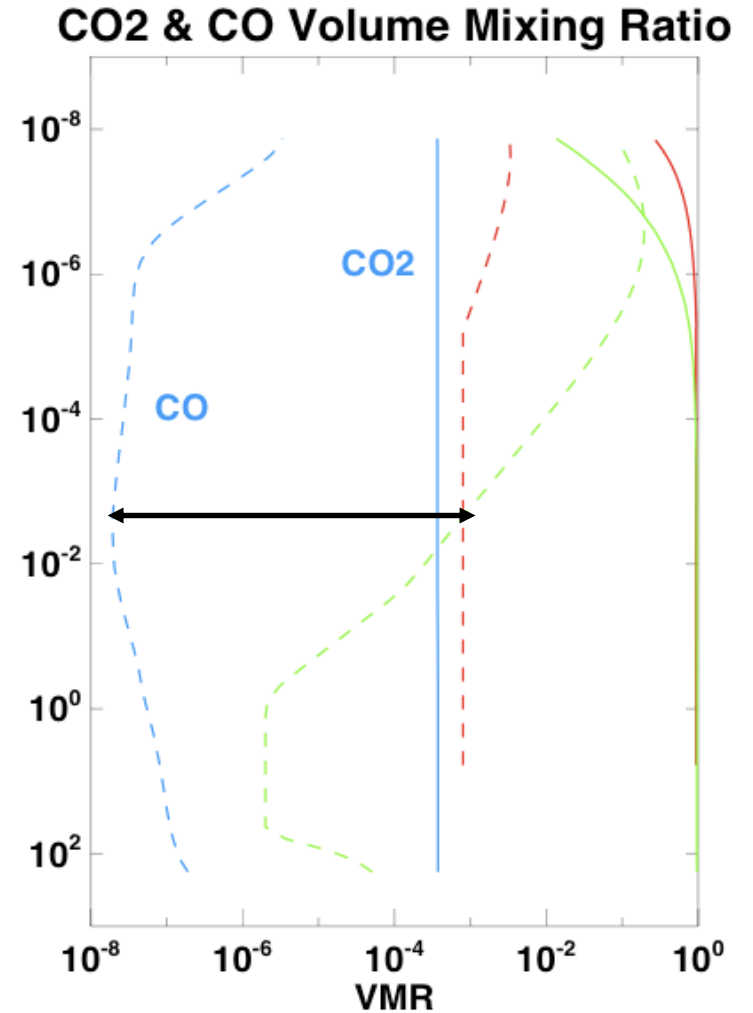
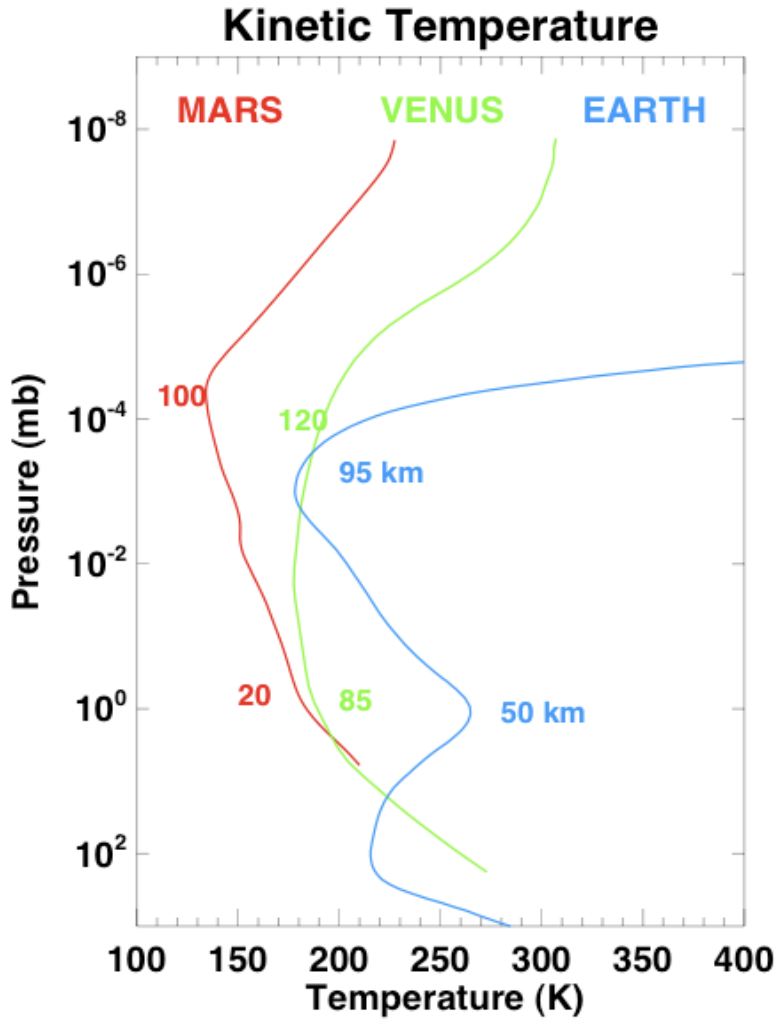
Funke et al. 2006

Coadded daytime spectra (red)

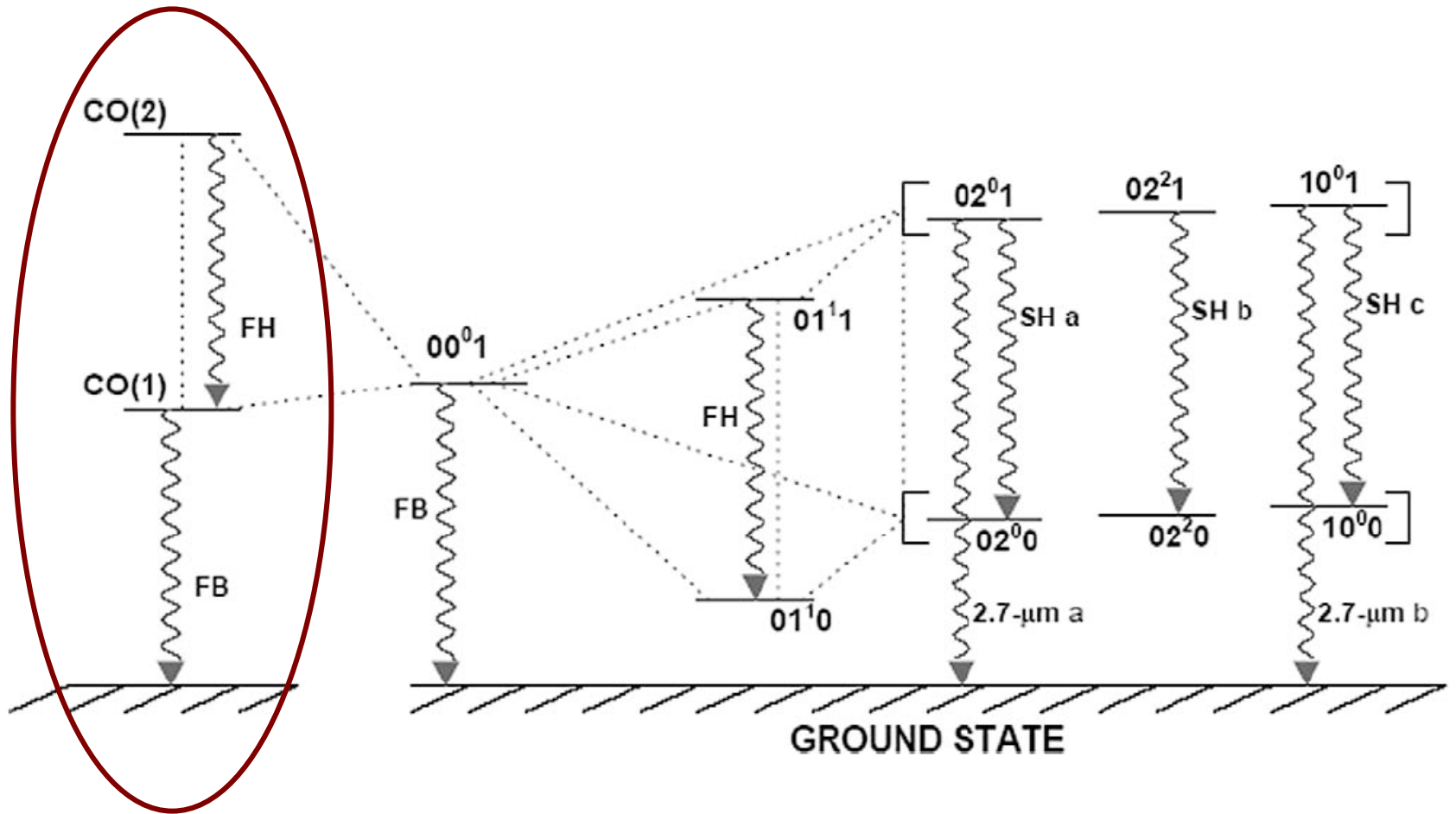
Coadded nighttime spectra (blue)

@ 68 km

Terrestrial Planets thermospheres



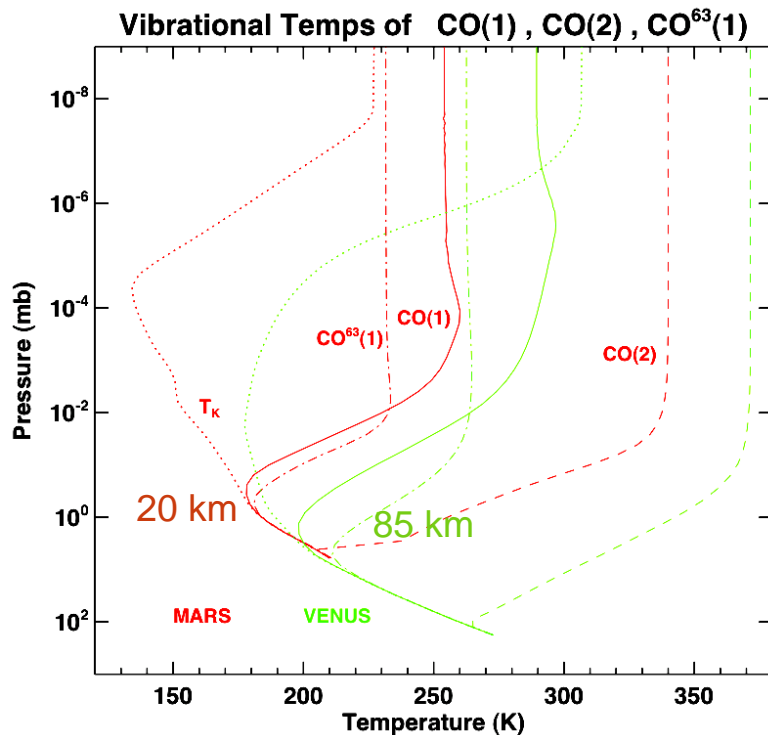
CO and CO₂ vibrational levels included in the non-LTE models (see poster by López-Valverde, this conference)



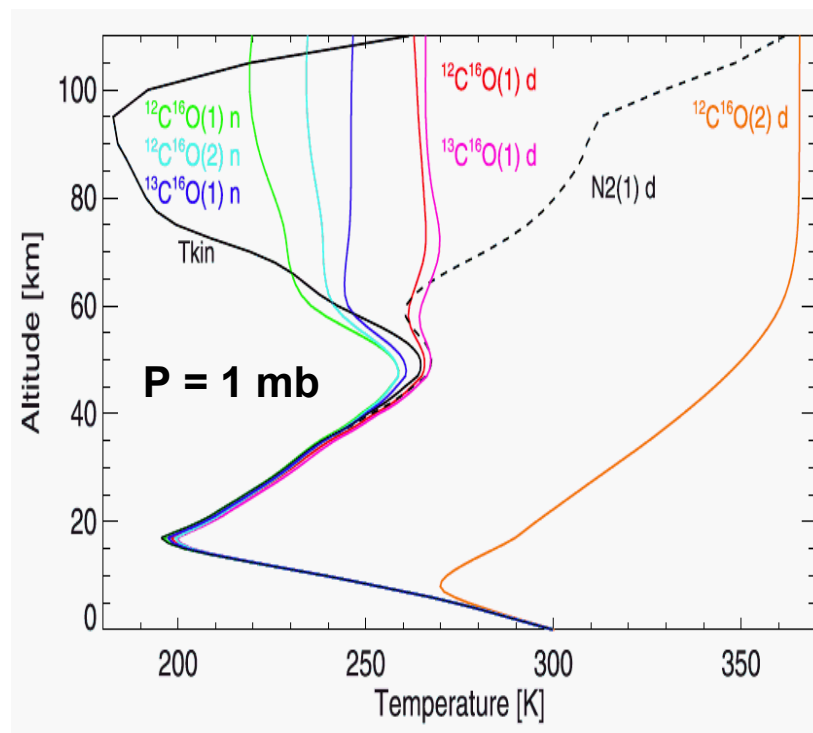
Non-LTE models results

CO Vibrational temperatures

MARS & VENUS



EARTH



Funke et al. 2006

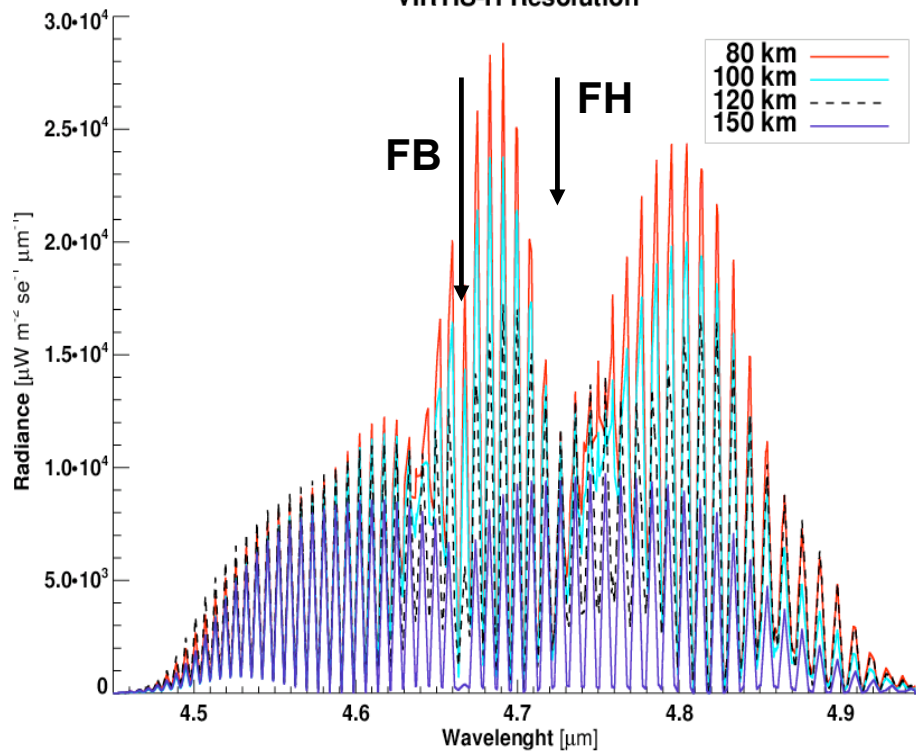
Non-LTE model results

Simulated CO spectra

RFM forward model: line-by-line calculations

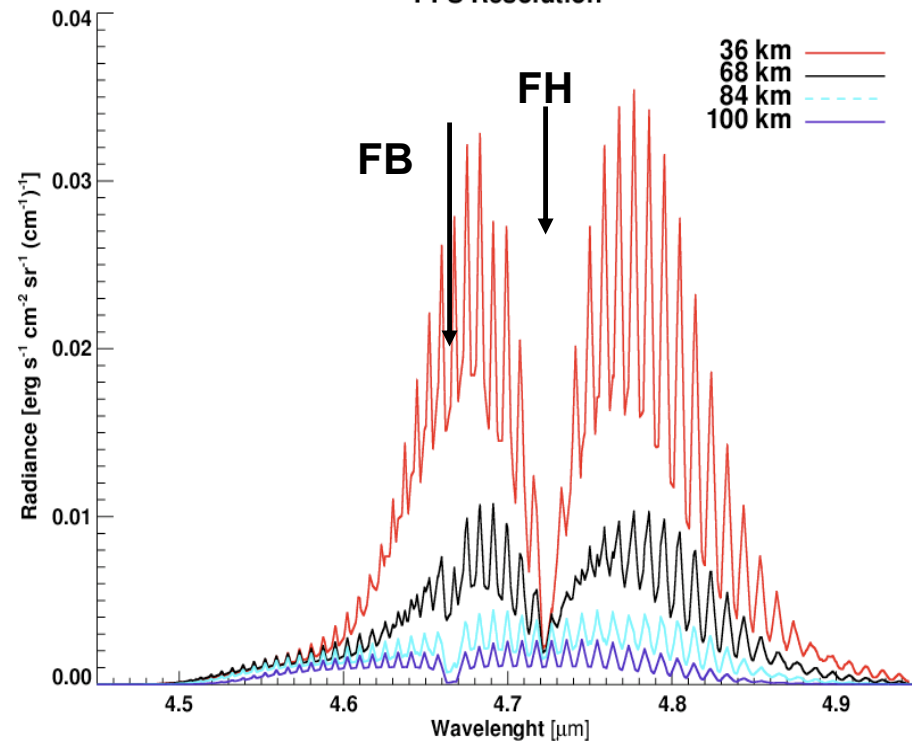
VENUS

VIRTIS-H Resolution



MARS

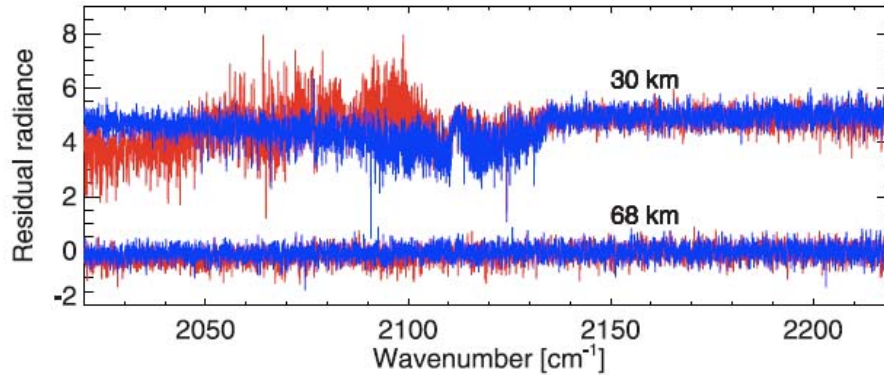
PFS Resolution



Comparison non-LTE models-data

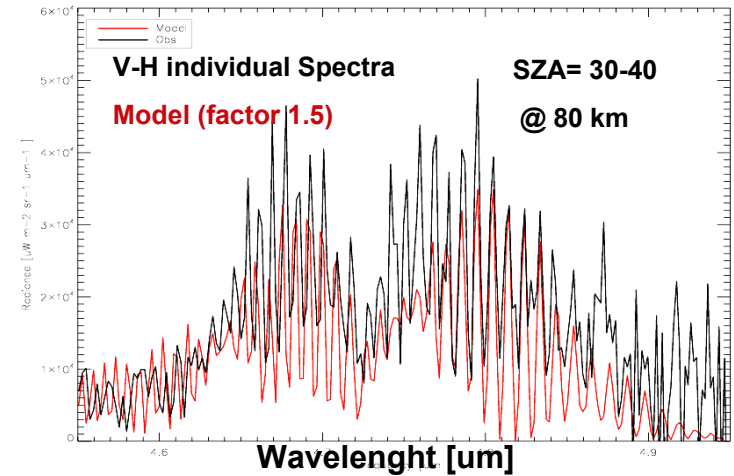
EARTH

Funke et al. 2006

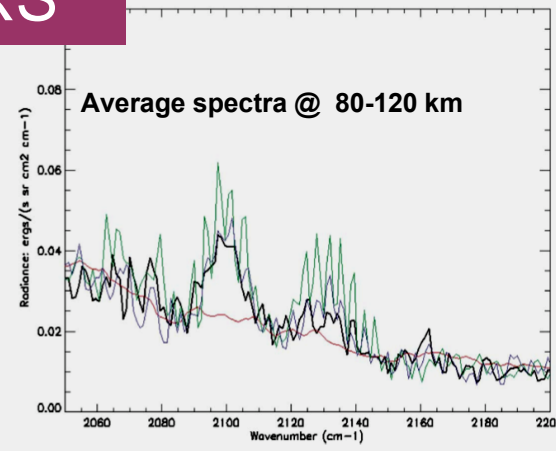
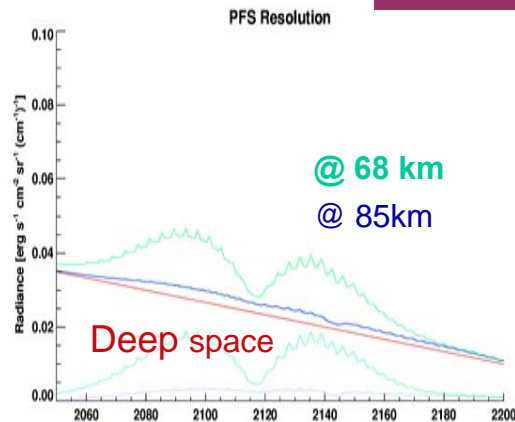


Simulated - measured coadded daytime/nighttime radiances (red/blue)

VENUS



MARS



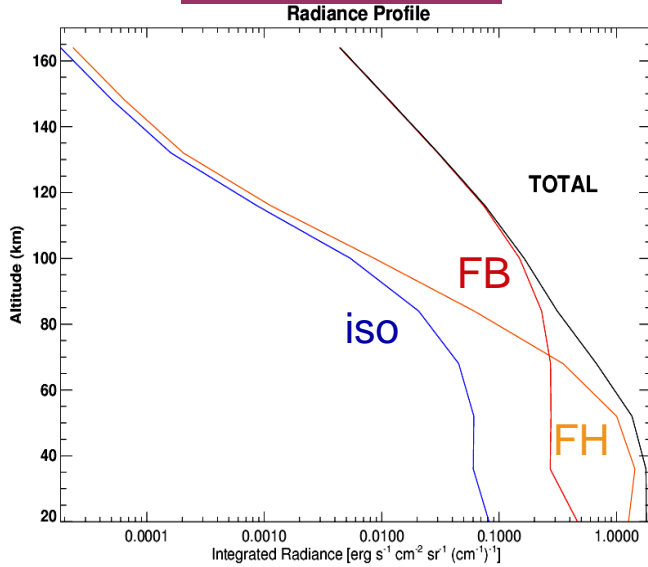
Formisano et al. 2006

Non-LTE model results

Simulated radiance profiles

integrated in all spectral region of 4.7- μm CO emission

MARS



**Optically thin
condition**

$$R \propto n^*$$

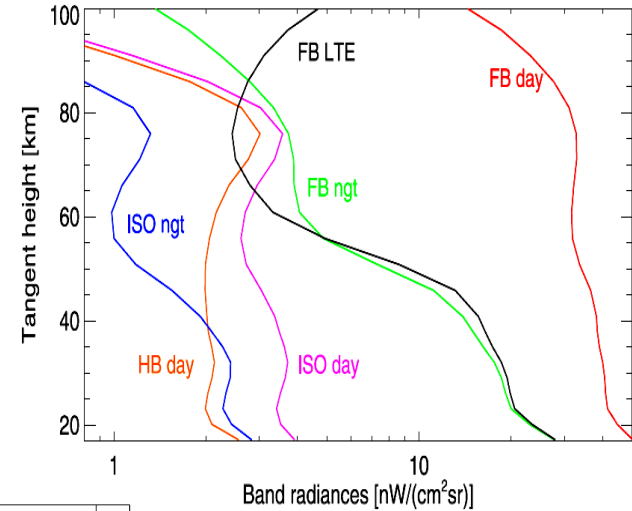
$$FB n^* > FH n^*$$

**Optically thick
condition**

$$R \propto J \propto T_{\text{vib}}$$

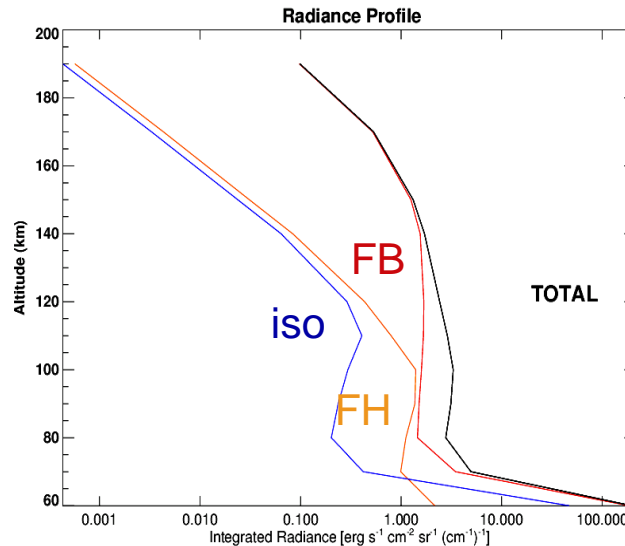
$$FH T_{\text{vib}} \gg FB T_{\text{vib}}$$

EARTH



Funke et al. 2006

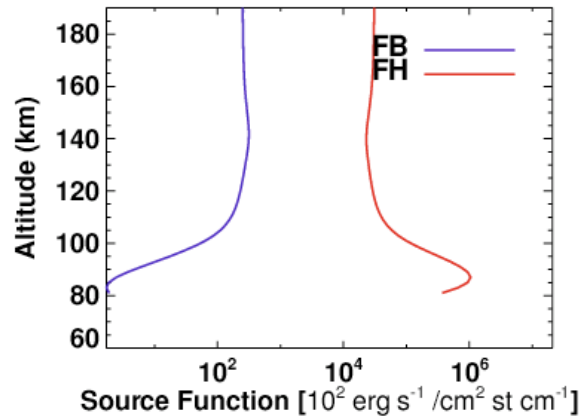
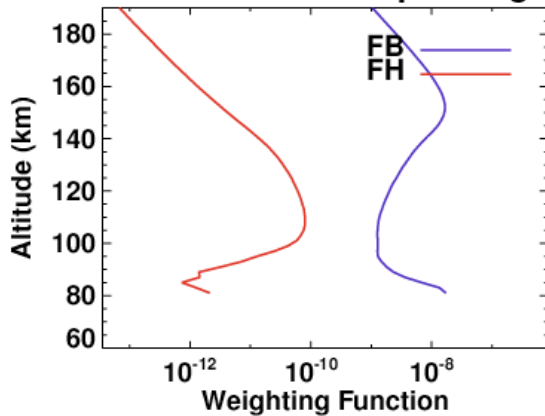
VENUS



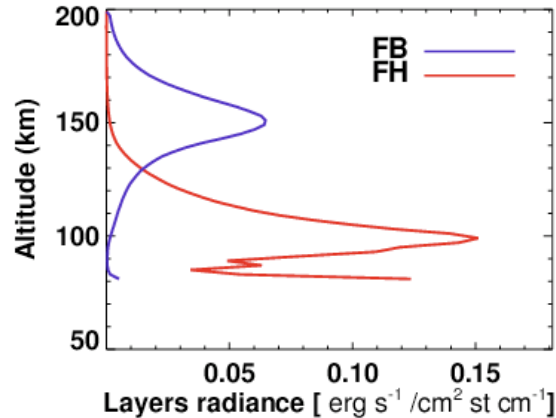
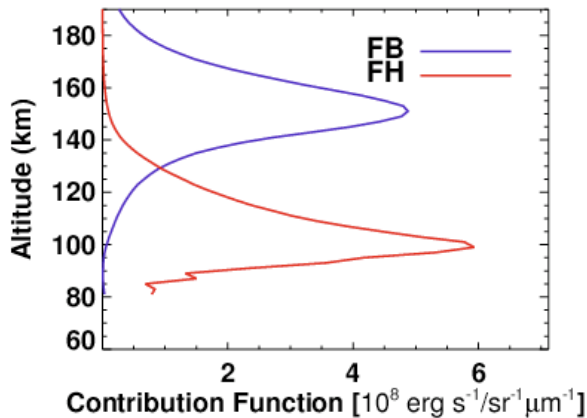
Contribution Function study for Venus

Spectrum convolved to VIRTIS-H resolution

Limb pointing @ 80 km

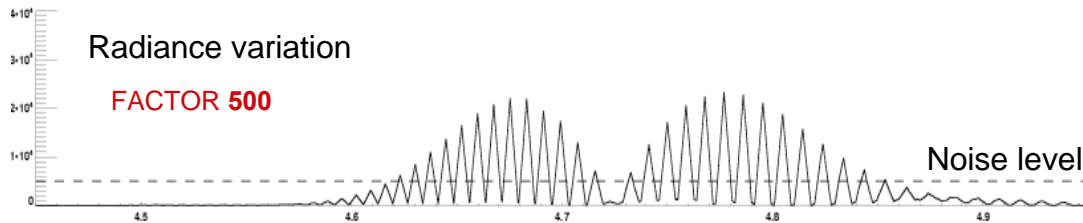
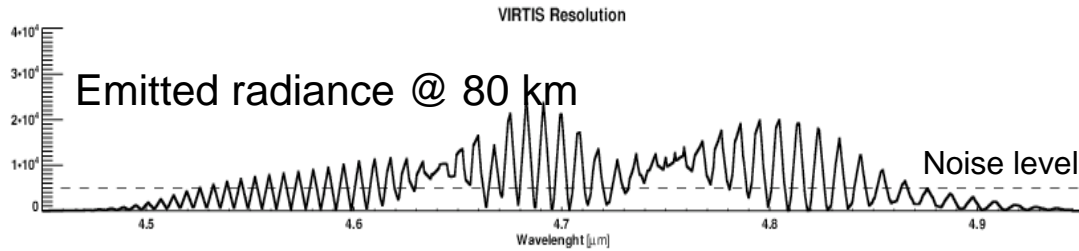


FB band is **optically thick**: its emission comes from the upper layers



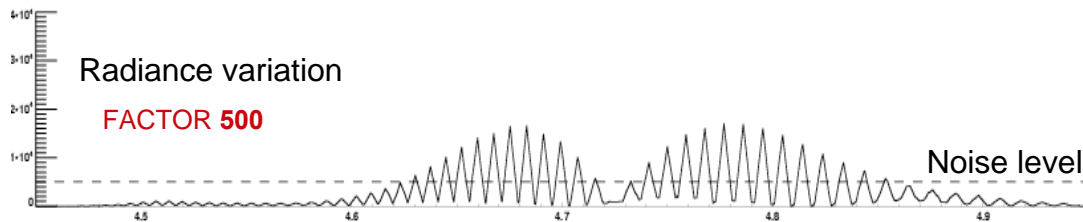
FH band is **optically thin** from above 100 km

Sensitivity study for VIRTIS measurements

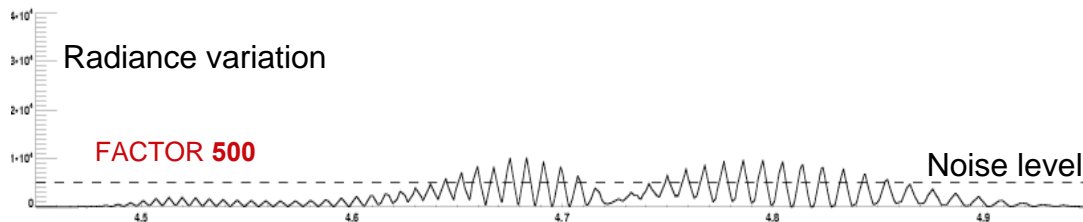


1% density perturbation @100 km

Density retrieved with uncertainty $\approx 125\%$

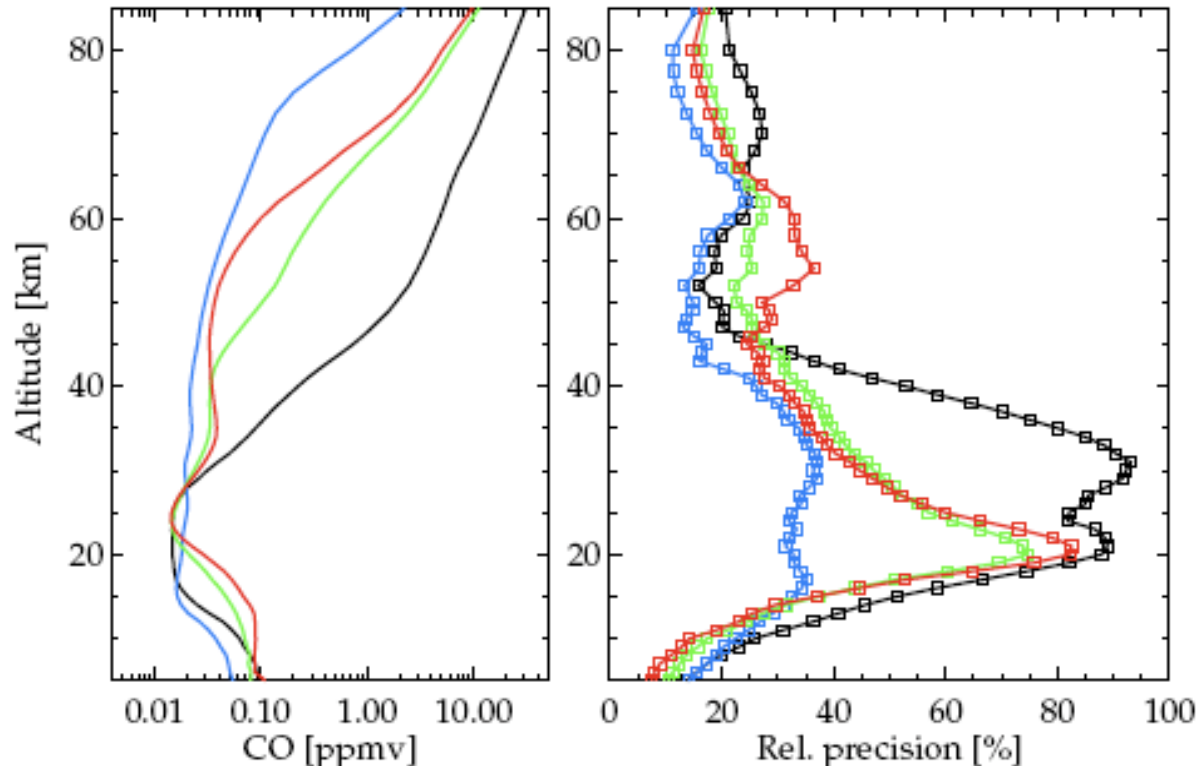


1% density perturbation @110 km



1% density perturbation @120 km

CO abundances on Earth from non-LTE retrievals of MIPAS




Funke et al. 2009

Retrieved average CO vmr profiles and relative precision:

Polar winter, Polar summer,

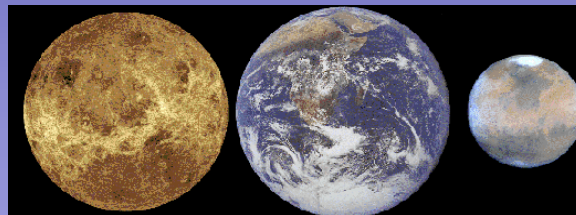
mid-latitudes, tropics

Summary & conclusions

1. Typical double ro-vibrational structure of CO diatomic molecule observed
2. On Mars & Venus FH band dominates CO limb spectra
3. On Earth lower CO VMR  bands are optically thinner
4. Good qualitative agreement with our non-LTE model simulations
5. High spectral resolution & good sensitivity: key to derive atmospheric parameter from measurements.
6. Nowadays, investigation on Earth is leader
7. VIRTIS is potentially useful to retrieve densities in the upper atmosphere of Venus using a non-LTE retrieval.

Thank you!

Good lunch!



VIRTIS-H Venus observations:

Orbit 712

Cube 2

